

Stationary Dyonic Regular and Black Hole Solutions

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We consider globally regular and black hole solutions in SU(2) Einstein-Yang-Mills-Higgs theory, coupled to a dilaton field. The basic solutions represent magnetic monopoles, monopole-antimonopole systems or black holes with monopole or dipole hair. When the globally regular solutions carry additionally electric charge, an angular momentum density results, except in the simplest spherically symmetric case. We evaluate the global charges of the solutions and their effective action, and analyze their dependence on the gravitational coupling strength. We show, that in the presence of a dilaton field, the black hole solutions satisfy a generalized Smarr type mass formula.

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I. INTRODUCTION

In Einstein-Maxwell (EM) theory the Kerr-Newman (KN) solutions represent stationary asymptotically flat black holes, characterized uniquely by their global charges: their mass M , their angular momentum J , their electric charge Q , and their magnetic charge P [1, 2]. Following Wheeler this uniqueness theorem of EM theory is often expressed as “EM black holes have no hair”.

In many unified theories, including string theory, dilatons appear. When a dilaton is coupled to EM theory, this has profound consequences for the black hole solutions. Not only do charged static Einstein-Maxwell-dilaton (EMD) black hole solutions exist for arbitrarily small horizon size [3], but also the staticity theorem of EM theory [4] does not generalize to EMD theory for arbitrary dilaton coupling constant γ : at the Kaluza-Klein value $\gamma_{\text{KK}} = \sqrt{3}$ stationary non-static black holes appear, whose horizon is non-rotating [5, 6], and beyond γ_{KK} even counterrotating black holes arise, whose horizon angular velocity and global angular momentum have opposite sign [6].

The EM uniqueness theorem, on the other hand, does not readily generalize to theories with non-Abelian gauge fields coupled to gravity [7]. The hairy black hole solutions of SU(2) Einstein-Yang-Mills (EYM) and Einstein-Yang-Mills-Higgs (EYMH) theory possess non-trivial magnetic fields outside their regular event horizon and are not uniquely characterized by their mass, their angular momentum, their electric and magnetic charge [8, 9, 10, 11, 12, 13, 14]. Furthermore, black hole solutions arise, which are static and not spherically symmetric, showing that Israel’s theorem [1, 2] does not generalize to non-Abelian theories, either [10, 11].

The coupling to non-Abelian fields not only gives rise to new types of black hole solutions, but also allows for globally regular solutions, not present in EM theory either [7, 9, 15]. These are stationary solutions with a spatially localized energy density of the matter fields and a finite mass, and are referred to as solitons when they are stable, and sphalerons when they possess unstable modes. The known globally regular solutions of EYM theory represent sphalerons, whereas the globally regular magnetic monopoles of EYMH theory are solitons, whose topological charge is proportional to their magnetic charge [16]. Besides magnetic monopoles EYMH theory contains a plethora of further globally regular solutions, representing for instance monopole-antimonopole pairs, chains, and vortex ring solutions [17].

It is an interesting question whether such globally regular solutions can be endowed with rotation, like their black hole counterparts can. When EYM black holes start to rotate, the time component of their gauge potential is excited, as expected. Surprisingly, however, not only a magnetic moment is induced by the rotation but also an electric charge [12, 18], and this seems to preclude the existence of globally regular rotating EYM sphalerons [19, 20].

Globally regular EYMH solutions, on the other hand, can carry electric charge, and the presence of a time component of the gauge potential renders the solutions stationary. Together the electric and magnetic fields then give rise to an angular momentum density, except in the spherically symmetric case [20, 21, 22]. Still, globally regular EYMH solutions with a non-vanishing global magnetic charge cannot rotate: their angular momentum vanishes [20, 21, 22]. But globally regular EYMH solutions with no global magnetic charge do possess a finite angular momentum. In fact, it is proportional to their electric charge [20], giving rise to a quantization condition for the angular momentum,

$$J = nQ(1 - \varepsilon) , \quad P = n\varepsilon , \quad \varepsilon = \frac{1}{2} [1 - (-1)^m] , \quad (1)$$

where m and n are two integers, characterizing the EYMH solutions [17].

Here we derive a mass formula for the stationary globally regular EYMH solutions in the presence of a dilaton. Then we address the dependence of the global charges and of the effective action of these solutions on the gravitational coupling strength. For a given type of solution, typically two branches of solutions arise, which bifurcate at a maximal value of the coupling, α_{\max} . For static solutions, the mass M exhibits a “spike” at α_{\max} [9, 17], since there the two branches must possess the same tangent w.r.t. α [23]. When stationary and rotating solutions are considered, in contrast, the mass branches may exhibit a “loop”, when considered as a function of α [22]. Here we show that for stationary and rotating solutions it is the effective action S^{eff} which may only exhibit a “spike” in the vicinity of the maximal value of the gravitational coupling constant. We illustrate this qualitative different behaviour of the mass M and the effective action S^{eff} for several sets of numerically constructed stationary and rotating solutions.

Turning to black holes again, we recall, that EM black holes satisfy the laws of black hole mechanics [24] and the Smarr mass formula [25]

$$M = 2TS + 2\Omega J + \tilde{\psi}_{\text{el}}Q + \tilde{\psi}_{\text{mag}}P , \quad (2)$$

where T represents the temperature of the black holes and S their entropy, Ω denotes their horizon angular velocity, and $\tilde{\psi}_{\text{el}}$ and $\tilde{\psi}_{\text{mag}}$ represent their horizon electric and magnetic potential, respectively.

In the presence of a dilaton an equivalent mass formula for EMD black holes is [14]

$$M = 2TS + 2\Omega J + \frac{D}{\gamma} + 2\tilde{\psi}_{\text{el}}Q , \quad (3)$$

with dilaton charge D and dilaton coupling constant γ . Interestingly, this second form of the mass formula also holds for the known non-Abelian black hole solutions of Einstein-Yang-Mills-dilaton (EYMD) theory [14].

Here we address stationary black holes of EYMH and Einstein-Yang-Mills-Higgs-dilaton (EYMHD) theory [26]. For these black holes the zeroth law of black hole mechanics holds [13], as well as a generalized first law [27]. We derive a mass formula for EYMHD black holes, based on the asymptotic expansion of the metric and the matter fields. The analytical mass formula represents a good criterion for the quality of numerically constructed EYMHD black hole solutions, also presented.

In section II we recall the SU(2) EYMHD action and the equations of motion. We discuss the stationary ansatz for the metric, the gauge potential, the Higgs field and the dilaton field, and we present the boundary conditions for globally regular and black hole solutions. In section III we address the physical properties of the solutions. We present the asymptotic expansion at infinity and the expansion at the horizon, needed to obtain the global charges and the horizon properties of the solutions. We evaluate the mass, the angular momentum and the effective action of the globally regular solutions in section IV, and discuss the dependence of these quantities on the coupling constant α . We illustrate these results for a set of numerically constructed solutions. We then derive the mass formula for the stationary black hole solutions in section V, presenting also numerical results. In section VI we present our conclusions.

II. EINSTEIN-YANG-MILLS-HIGGS-DILATON SOLUTIONS

After recalling the SU(2) EYMHD action and the general set of equations of motion, we discuss the ansatz for the stationary non-Abelian globally regular and black hole solutions. The ansatz for the metric represents the stationary axially symmetric Lewis-Papapetrou metric in isotropic coordinates. The ansatz for the gauge potential and the Higgs field includes two integers, m and n , related to the polar and azimuthal angles. For monopole-antimonopole chains the integer m counts the total number of poles on the symmetry axis, while the integer n gives the magnitude of the magnetic charge of each pole. As implied by the boundary conditions, the stationary axially symmetric solutions are asymptotically flat, and the black hole solutions possess a regular event horizon.

A. SU(2) EYMHD Action

We consider the SU(2) Einstein-Yang-Mills-dilaton action

$$\mathcal{S} = \int \left(\frac{R}{16\pi G} + L_M \right) \sqrt{-g} d^4x , \quad (4)$$

where R is the scalar curvature, and the matter Lagrangian L_M is given by

$$L_M = -\frac{1}{2}\partial_\mu\Psi\partial^\mu\Psi - \frac{1}{2}e^{2\kappa\Psi}\text{Tr}(F_{\mu\nu}F^{\mu\nu}) - \frac{1}{4}\text{Tr}(D_\mu\Phi D^\mu\Phi) - \frac{\lambda}{8}e^{-2\kappa\Psi}\text{Tr}(\Phi^2 - v^2)^2, \quad (5)$$

with dilaton field Ψ , gauge field strength tensor $F_{\mu\nu} = \partial_\mu A_\nu - \partial_\nu A_\mu + ie[A_\mu, A_\nu]$, gauge field $A_\mu = A_\mu^a \tau_a/2$, Higgs field in the adjoint representation $\Phi = \tau^a \Phi^a$, gauge covariant derivative $D_\mu = \nabla_\mu + ie[A_\mu, \cdot]$, and Newton's constant G , dilaton coupling constant κ , Yang-Mills coupling constant e , Higgs self-coupling constant λ , and Higgs vacuum expectation value v .

The nonzero vacuum expectation value of the Higgs field breaks the non-Abelian $SU(2)$ gauge symmetry to the Abelian $U(1)$ symmetry. The particle spectrum of the theory then consists of a massless photon, two massive vector bosons of mass $M_W = ev$, and a massive Higgs field $M_H = \sqrt{2\lambda}v$. In the limit $\lambda = 0$ the Higgs field also becomes massless. The dilaton is massless as well.

Including a boundary term [28], variation of the action with respect to the metric and the matter fields leads, respectively, to the Einstein equations

$$G_{\mu\nu} = R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R = 8\pi GT_{\mu\nu} \quad (6)$$

with stress-energy tensor

$$\begin{aligned} T_{\mu\nu} &= g_{\mu\nu}L_M - 2\frac{\partial L_M}{\partial g^{\mu\nu}} \\ &= \partial_\mu\Psi\partial_\nu\Psi - \frac{1}{2}g_{\mu\nu}\partial_\alpha\Psi\partial^\alpha\Psi + 2e^{2\kappa\Psi}\text{Tr}(F_{\mu\alpha}F_{\nu\beta}g^{\alpha\beta} - \frac{1}{4}g_{\mu\nu}F_{\alpha\beta}F^{\alpha\beta}) \\ &\quad + \frac{1}{2}\text{Tr}\left(D_\mu\Phi D_\nu\Phi - \frac{1}{2}g_{\mu\nu}D_\alpha\Phi D^\alpha\Phi\right) - \frac{\lambda}{8}g_{\mu\nu}e^{-2\kappa\Psi}\text{Tr}(\Phi^2 - v^2)^2, \end{aligned} \quad (7)$$

and the matter field equations,

$$D_\mu(e^{2\kappa\Psi}F^{\mu\nu}) = \frac{1}{4}ie[\Phi, D^\nu\Phi], \quad (8)$$

$$\square\Psi = \kappa e^{2\kappa\Psi}\text{Tr}(F_{\mu\nu}F^{\mu\nu}) - \frac{\lambda}{4}\kappa e^{-2\kappa\Psi}\text{Tr}(\Phi^2 - v^2)^2, \quad (9)$$

where $\square\Psi = \Psi_{;\sigma}{}^{;\sigma}$, and

$$D_\mu D^\mu\Phi = \lambda e^{-2\kappa\Psi}\text{Tr}(\Phi^2 - v^2)\Phi. \quad (10)$$

B. Stationary Axially Symmetric Ansatz

The system of partial differential equations, Eq. (6), Eq. (9), Eq. (8), and Eq. (10) is highly non-linear and complicated. In order to generate solutions to these equations, one profits from the use of symmetries, simplifying the equations.

Here we consider solutions, which are both stationary and axially symmetric. We therefore impose on the spacetime the presence of two commuting Killing vector fields, ξ (asymptotically timelike) and η (asymptotically spacelike). Since the Killing vector fields commute, we may adopt a system of adapted coordinates, say $\{t, r, \theta, \varphi\}$, such that

$$\xi = \partial_t, \quad \eta = \partial_\varphi. \quad (11)$$

In these coordinates the metric is independent of t and φ . We also assume that the symmetry axis of the spacetime, the set of points where $\eta = 0$, is regular, and satisfies the elementary flatness condition

$$\frac{X_{;\mu}X^{;\mu}}{4X} = 1, \quad X = \eta^\mu\eta_\mu. \quad (12)$$

Apart from the symmetry requirement on the metric ($\mathcal{L}_\xi g = \mathcal{L}_\eta g = 0$, i.e., $g_{\mu\nu} = g_{\mu\nu}(r, \theta)$), we impose that the matter fields are also symmetric under the spacetime transformations generated by ξ and η .

This implies for the dilaton field

$$\mathcal{L}_\xi \Psi = \mathcal{L}_\eta \Psi = 0 , \quad (13)$$

so Ψ depends on r and θ only. Introducing two compensating $\text{su}(2)$ -valued functions W_ξ and W_η , the concept of generalised symmetry [20, 29] requires for the Higgs field

$$\mathcal{L}_\xi \Phi = ie[\Phi, W_\xi] , \quad \mathcal{L}_\eta \Phi = ie[\Phi, W_\eta] , \quad (14)$$

and for the gauge potential $A = A_\mu dx^\mu$,

$$\begin{aligned} (\mathcal{L}_\xi A)_\mu &= D_\mu W_\xi , \\ (\mathcal{L}_\eta A)_\mu &= D_\mu W_\eta , \end{aligned} \quad (15)$$

where W_ξ and W_η satisfy

$$\mathcal{L}_\xi W_\eta - \mathcal{L}_\eta W_\xi + ie[W_\xi, W_\eta] = 0 . \quad (16)$$

Performing a gauge transformation to set $W_\xi = 0$, leaves Φ , A and W_η independent of t .

By virtue of the Frobenius condition and the circularity theorem, the metric can then be written in the Lewis-Papapetrou form, which in isotropic coordinates reads

$$ds^2 = -f dt^2 + \frac{m}{f} [dr^2 + r^2 d\theta^2] + \sin^2 \theta r^2 \frac{l}{f} \left[d\varphi - \frac{\omega}{r} dt \right]^2 , \quad (17)$$

where f , m , l and ω are functions of r and θ only.

The z -axis represents the symmetry axis. The regularity condition along the z -axis Eq. (12) requires

$$m|_{\theta=0,\pi} = l|_{\theta=0,\pi} . \quad (18)$$

The event horizon of stationary black hole solutions resides at a surface of constant radial coordinate, $r = r_H$, and is characterized by the condition $f(r_H, \theta) = 0$ [12]. The Killing vector field

$$\chi = \xi + \frac{\omega_H}{r_H} \eta , \quad (19)$$

is orthogonal to and null on the horizon [24]. The ergosphere, defined as the region in which $\xi_\mu \xi^\mu$ is positive, is bounded by the event horizon and by the surface where

$$-f + \sin^2 \theta \frac{l}{f} \omega^2 = 0 . \quad (20)$$

For the gauge fields we employ a generalized ansatz [12, 13, 22], which trivially fulfils both the symmetry constraints Eq. (15) and Eq. (16) and the circularity conditions,

$$A_\mu dx^\mu = \left(B_1 \frac{\tau_r^{(n,m)}}{2e} + B_2 \frac{\tau_\theta^{(n,m)}}{2e} \right) dt + A_\varphi \left(d\varphi - \frac{\omega}{r} dt \right) + \left(\frac{H_1}{r} dr + (1 - H_2) d\theta \right) \frac{\tau_\varphi^{(n)}}{2e} , \quad (21)$$

$$A_\varphi = -n \sin \theta \left(H_3 \frac{\tau_r^{(n,m)}}{2e} + (1 - H_4) \frac{\tau_\theta^{(n,m)}}{2e} \right) , \quad (22)$$

and the appropriate Ansatz for the Higgs field is then given by [13, 22]

$$\Phi = v \left(\Phi_1 \tau_r^{(n,m)} + \Phi_2 \tau_\theta^{(n,m)} \right) , \quad (23)$$

where n and m are integers. The symbols $\tau_r^{(n,m)}$, $\tau_\theta^{(n,m)}$ and $\tau_\varphi^{(n)}$ denote the dot products of the Cartesian vector of Pauli matrices, $\vec{\tau} = (\tau_x, \tau_y, \tau_z)$, with the spatial unit vectors

$$\begin{aligned} \hat{e}_r^{(n,m)} &= (\sin(m\theta) \cos(n\varphi), \sin(m\theta) \sin(n\varphi), \cos(m\theta)) , \\ \hat{e}_\theta^{(n,m)} &= (\cos(m\theta) \cos(n\varphi), \cos(m\theta) \sin(n\varphi), -\sin(m\theta)) , \\ \hat{e}_\varphi^{(n)} &= (-\sin(n\varphi), \cos(n\varphi), 0) , \end{aligned} \quad (24)$$

respectively. Like the dilaton field function Ψ , the gauge field functions B_i and H_i and the Higgs field functions Φ_i depend only on the coordinates r and θ .

The ansatz is form-invariant under Abelian gauge transformations U [10, 30]

$$U = \exp \left(\frac{i}{2} \tau_\varphi^{(n)} \Gamma(r, \theta) \right) . \quad (25)$$

With respect to this residual gauge degree of freedom we choose the gauge fixing condition $r\partial_r H_1 - \partial_\theta H_2 = 0$. For the gauge field ansatz, Eqs. (21)-(22), the compensating matrix W_η is given by

$$W_\eta = n \frac{\tau_z}{2e} . \quad (26)$$

C. Dimensionless Quantities

Let us now introduce the dimensionless quantities, beginning with the dimensionless coupling constants α , β and γ

$$v = \frac{\alpha}{\sqrt{4\pi G}} , \quad \lambda = e^2 \beta^2 , \quad \kappa = \frac{\sqrt{4\pi G}}{\alpha} \gamma . \quad (27)$$

The dimensionless coordinate x is given by

$$r = \frac{\sqrt{4\pi G}}{e\alpha} x , \quad (28)$$

the dimensionless electric gauge field functions \bar{B}_1 and \bar{B}_2 are

$$B_1 = \frac{e\alpha}{\sqrt{4\pi G}} \bar{B}_1 , \quad B_2 = \frac{e\alpha}{\sqrt{4\pi G}} \bar{B}_2 , \quad (29)$$

and the dimensionless dilaton function ψ is

$$\Psi = \frac{\alpha}{\sqrt{4\pi G}} \psi . \quad (30)$$

Introducing these dimensionless quantities into the EOMs, the resulting equations depend only on the parameters α , β , and γ . Note, that in the limit $\gamma \rightarrow 0$ the dilaton decouples and the equations of EYM theory are obtained.

D. Boundary Conditions

Boundary conditions at infinity

To obtain asymptotically flat solutions, we impose on the metric functions the boundary conditions at infinity

$$f|_{x=\infty} = m|_{x=\infty} = l|_{x=\infty} = 1 , \quad \omega|_{x=\infty} = 0 . \quad (31)$$

For the dilaton function we choose

$$\psi|_{x=\infty} = 0 , \quad (32)$$

since any finite value of the dilaton field at infinity can always be transformed to zero via $\psi \rightarrow \psi - \psi(\infty)$, $x \rightarrow xe^{-\gamma\psi(\infty)}$.

The asymptotic values of the Higgs field functions Φ_i are

$$\Phi_1|_{x=\infty} = 1 , \quad \Phi_2|_{x=\infty} = 0 . \quad (33)$$

We further impose, that the two electric gauge field functions \bar{B}_i satisfy

$$\bar{B}_1|_{x=\infty} = \nu , \quad \bar{B}_2|_{x=\infty} = 0 , \quad (34)$$

where the asymptotic value ν is restricted to $0 \leq |\nu| < 1$, and that the magnetic gauge field functions H_i satisfy

$$H_1|_{x=\infty} = 0 , \quad H_2|_{x=\infty} = 1 - m , \quad (35)$$

$$H_3|_{x=\infty} = \frac{\cos \theta - \cos(m\theta)}{\sin \theta} \quad m \text{ odd} , \quad H_3|_{x=\infty} = \frac{1 - \cos(m\theta)}{\sin \theta} \quad m \text{ even} , \quad (36)$$

$$H_4|_{x=\infty} = 1 - \frac{\sin(m\theta)}{\sin \theta} . \quad (37)$$

Boundary conditions at the origin

To obtain globally regular solutions, we must impose appropriate boundary conditions at the origin. Regularity requires for the metric functions the boundary conditions

$$\partial_x f|_{x=0} = \partial_x m|_{x=0} = \partial_x l|_{x=0} = 0 , \quad \omega|_{x=0} = 0 , \quad (38)$$

and for the dilaton function

$$\partial_x \psi|_{x=0} = 0 , \quad (39)$$

the gauge field functions H_i satisfy

$$H_1|_{x=0} = H_3|_{x=0} = 0 , \quad H_2|_{x=0} = H_4|_{x=0} = 1 , \quad (40)$$

while for even m the gauge field functions \bar{B}_i and the Higgs functions Φ_i satisfy

$$[\sin(m\theta)\Phi_1 + \cos(m\theta)\Phi_2]|_{x=0} = 0 , \quad (41)$$

$$\partial_x [\cos(m\theta)\Phi_1 - \sin(m\theta)\Phi_2]|_{x=0} = 0 , \quad (42)$$

$$[\sin(m\theta)\bar{B}_1 + \cos(m\theta)\bar{B}_2 = 0]|_{x=0} = 0 , \quad (43)$$

$$\partial_x [\cos(m\theta)\bar{B}_1 - \sin(m\theta)\bar{B}_2]|_{x=0} = 0 , \quad (44)$$

whereas for odd m they satisfy $\bar{B}_i|_{x=0} = \Phi_i|_{x=0} = 0$.

Boundary conditions at the horizon

The event horizon of stationary black hole solutions resides at a surface of constant radial coordinate, $x = x_H$, and is characterized by the condition $f(x_H, \theta) = 0$ [12].

Regularity at the horizon then requires the following boundary conditions for the metric functions

$$f|_{x=x_H} = m|_{x=x_H} = l|_{x=x_H} = 0 , \quad \omega|_{x=x_H} = \omega_H = \text{const.} , \quad (45)$$

for the dilaton function

$$\partial_x \psi|_{x=x_H} = 0 , \quad (46)$$

while the Higgs and the magnetic gauge field functions satisfy

$$\partial_x \Phi_1|_{x=x_H} = \partial_x \Phi_2|_{x=x_H} = 0 , \quad (47)$$

$$H_1|_{x=x_H} = 0 , \quad \partial_x H_2|_{x=x_H} = \partial_x H_3|_{x=x_H} = \partial_x H_4|_{x=x_H} = 0 , \quad (48)$$

with the gauge condition $\partial_\theta H_1 = 0$ taken into account [12]. The boundary conditions for the electric gauge field functions are obtained from the requirement that for non-Abelian solutions the electrostatic potential is constant at the horizon [13]

$$\tilde{\Psi}_{\text{el}} \frac{\tau_z}{2} = -\chi^\mu A_\mu|_{r=r_H} . \quad (49)$$

Defining the dimensionless electrostatic potential $\tilde{\psi}_{\text{el}}$,

$$\tilde{\psi}_{\text{el}} = \frac{\sqrt{4\pi G}}{\alpha} \tilde{\Psi}_{\text{el}} , \quad (50)$$

and the dimensionless horizon angular velocity Ω ,

$$\Omega = \frac{\omega_H}{x_H} , \quad (51)$$

yields the boundary conditions

$$\bar{B}_1|_{x=x_H} = n\Omega \cos m\theta , \quad \bar{B}_2|_{x=x_H} = -n\Omega \sin m\theta . \quad (52)$$

Boundary conditions along the symmetry axis

The boundary conditions along the z -axis ($\theta = 0$ and $\theta = \pi$) are determined by the symmetries. For the positive z -axis they are given by

$$\partial_\theta f|_{\theta=0} = \partial_\theta m|_{\theta=0} = \partial_\theta l|_{\theta=0} = \partial_\theta \omega|_{\theta=0} = 0 , \quad (53)$$

$$\partial_\theta \psi|_{\theta=0} = 0 , \quad (54)$$

$$H_1|_{\theta=0} = H_3|_{\theta=0} = 0 , \quad \partial_\theta H_2|_{\theta=0} = \partial_\theta H_4|_{\theta=0} = 0 , \quad (55)$$

$$\bar{B}_2|_{\theta=0} = 0 , \quad \partial_\theta \bar{B}_1|_{\theta=0} = 0 , \quad (56)$$

$$\Phi_2|_{\theta=0} = 0 , \quad \partial_\theta \Phi_1|_{\theta=0} = 0 . \quad (57)$$

The analogous conditions hold on the negative z -axis. We note, that the globally regular solutions are symmetric w.r.t. the xy -plane. For the black hole solutions, this symmetry is broken via the boundary conditions of the time component of the gauge field.

In addition, regularity on the z -axis requires condition Eq. (18) for the metric functions to be satisfied, and regularity of the energy density on the z -axis requires

$$H_2|_{\theta=0} = H_4|_{\theta=0} . \quad (58)$$

III. PROPERTIES OF REGULAR AND BLACK HOLE SOLUTIONS

We derive the properties of the stationary axially symmetric solutions from the expansions of their metric and matter field functions at infinity, at the origin and at the horizon. The expansion at infinity yields the global charges of the solutions, the expansion at the horizon yields the horizon properties of the black holes.

A. Asymptotic Expansion

The asymptotic expansion depends on the integers m and n . Here we restrict to odd winding number n , since the analysis for the even case seems to be ‘prohibitively complicated’. We then obtain for $\beta = 0$

$$\begin{aligned}
H_1 &= -\frac{C_1 \sin \theta}{x} - \frac{C_3 \sin(2\theta)}{x^2} + O\left(\frac{1}{x^3}\right), \\
H_2 &= (1-m) - \frac{C_1 \cos \theta}{x} - \frac{C_3 \cos(2\theta)}{x^2} + O\left(\frac{1}{x^3}\right), \\
H_3 &= \frac{\cos(\varepsilon\theta) - \cos(m\theta)}{\sin \theta} + \frac{C_6 \sin \theta}{x} + O\left(\frac{1}{x^2}\right), \\
H_4 &= \left(1 - \frac{\sin(m\theta)}{\sin \theta}\right) - \frac{C_1 \cos(\varepsilon\theta)}{x} - \frac{C_3 \cos^2 \theta + C_1 C_6 \sin^2 \theta}{x^2} + O\left(\frac{1}{x^3}\right), \\
\bar{B}_1 &= \nu - \frac{Q}{x} + \frac{2Q(\mu - \gamma D) - \nu C_1^2 \sin^2 \theta + 2C_7 \cos \theta}{2x^2} + O\left(\frac{1}{x^3}\right), \\
\bar{B}_2 &= \frac{\nu C_1 \sin \theta}{x} - \frac{(C_1 Q - \nu C_3 \cos \theta) \sin \theta}{x^2} + O\left(\frac{1}{x^3}\right), \\
f &= 1 - \frac{2\mu}{x} + \frac{2\mu^2 + \alpha^2(Q^2 + P^2) + C_4 \cos \theta}{x^2} + O\left(\frac{1}{x^3}\right), \\
m &= 1 + \frac{C_5 \cos(2\theta) + [-\mu^2 + \alpha^2(Q^2 + P^2 - D^2 - C_2^2)] \sin^2 \theta}{x^2} + O\left(\frac{1}{x^3}\right), \\
l &= 1 + \frac{C_5}{x^2} + O\left(\frac{1}{x^3}\right), \\
\omega &= \frac{2\zeta}{x^2} + O\left(\frac{1}{x^3}\right), \\
\psi &= -\frac{D}{x} - \frac{\gamma(Q^2 - P^2) - 2C_8 \cos \theta}{2x^2} + O\left(\frac{1}{x^3}\right), \\
\Phi_1 &= 1 + \frac{C_2}{x} - \frac{C_1^2 \sin^2 \theta - 2C_9 \cos \theta}{2x^2} + O\left(\frac{1}{x^3}\right), \\
\Phi_2 &= \frac{C_1 \sin \theta}{x} + \frac{(C_1 C_2 + C_3 \cos \theta) \sin \theta}{x^2} + O\left(\frac{1}{x^3}\right),
\end{aligned} \tag{59}$$

where

$$\varepsilon = \frac{1}{2}[1 - (-1)^m], \quad P = n\varepsilon. \tag{60}$$

For generic $\beta \neq 0$ solutions, the expansions remain valid with $C_2 = C_9 = 0$. At first sight the power law decay of the Higgs field then appears surprising, since $\beta \neq 0$ renders the Higgs massive and should thus lead to an exponential decay. However, this power law decay represents a gauge artifact and can be removed by the gauge transformation

$$U = \exp(i\Gamma\tau_\varphi^n/2), \tag{61}$$

with

$$\Gamma = -(1-m)\theta + \frac{C_1 \sin \theta}{x} + \frac{C_3 \sin(2\theta)}{2x^2}. \tag{62}$$

Performing this gauge transformation leads to $\Phi = v\tau_r^{(n,1)} + O(1/r^3)$ (note that the integer m has been transformed away). We further obtain trivial gauge field functions H_1 and H_2 (up to order $O(1/x^3)$).

B. Global Charges

The expansion coefficients M , J , Q , P and D correspond to the global charges of the solutions. The dimensionless mass M and angular momentum J of the solutions are obtained from the asymptotic expansion of the metric

$$M = \frac{1}{2\alpha^2} \lim_{x \rightarrow \infty} x^2 \partial_x f = \frac{\mu}{\alpha^2} , \quad J = \frac{1}{2\alpha^2} \lim_{x \rightarrow \infty} x^2 \omega = \frac{\zeta}{\alpha^2} , \quad (63)$$

These correspond to the expressions obtained from the respective Komar integrals, as shown in sections IV and V for the globally regular and black hole solutions, respectively. Note that M is the mass in units of $4\pi v/e$, whereas μ corresponds to the mass in units of Planck mass. Likewise, the dimensionless electric charge Q and the dimensionless magnetic charge P are given by

$$Q = - \lim_{x \rightarrow \infty} x (\bar{B}_1 - \nu) , \quad P = \frac{n}{2} (1 - (-1)^m) , \quad (64)$$

while the dimensionless dilaton charge D is given by

$$D = \lim_{x \rightarrow \infty} x^2 \partial_x \psi . \quad (65)$$

C. Expansion at the Horizon

Expanding the metric and matter field functions at the horizon in powers of

$$\delta = \frac{x}{x_H} - 1 \quad (66)$$

yields to lowest order

$$\begin{aligned} H_1 &= \delta \left(1 - \frac{1}{2} \delta \right) H_{11} + O(\delta^3) , \\ H_2 &= H_{20} + O(\delta^2) , \\ H_3 &= H_{30} + O(\delta^2) , \\ H_4 &= H_{40} + O(\delta^2) , \\ \bar{B}_1 &= n \frac{\omega_H}{x_H} \cos(m\theta) + O(\delta^2) , \\ \bar{B}_2 &= -n \frac{\omega_H}{x_H} \sin(m\theta) + O(\delta^2) , \\ f &= \delta^2 f_2 (1 - \delta) + O(\delta^4) , \\ m &= \delta^2 m_2 (1 - 3\delta) + O(\delta^4) , \\ l &= \delta^2 l_2 (1 - 3\delta) + O(\delta^4) , \\ \omega &= \omega_H (1 + \delta) + O(\delta^2) , \\ \psi &= \psi_0 + O(\delta^2) , \\ \Phi_1 &= \Phi_{10} + O(\delta^2) , \\ \Phi_2 &= \Phi_{20} + O(\delta^2) . \end{aligned} \quad (67)$$

The expansion coefficients f_2 , m_2 , l_2 , ψ_0 , H_{11} , H_{20} , H_{30} , H_{40} , Φ_{10} , Φ_{20} are functions of the variable θ . Among these coefficients the following relations hold,

$$0 = \frac{\partial_\theta m_2}{m_2} - 2 \frac{\partial_\theta f_2}{f_2} , \quad (68)$$

$$H_{11} = \partial_\theta H_{20} . \quad (69)$$

D. Horizon Properties

With help of the above expansion we obtain the horizon properties of the SU(2) EYMHD black hole solutions. The first quantity of interest is the area of the black hole horizon. The dimensionless area A_H is given by

$$A_H = 2\pi \int_0^\pi d\theta \sin \theta \frac{\sqrt{l_2 m_2}}{f_2} x_H^2 , \quad (70)$$

and the dimensionless entropy S by

$$S = \frac{A_H}{4} . \quad (71)$$

The surface gravity of the black hole solutions is obtained from [24]

$$\kappa_{\text{sg}}^2 = -\frac{1}{2}(\nabla_\mu \chi_\nu)(\nabla^\mu \chi^\nu) , \quad (72)$$

with Killing vector χ , Eq. (19). Inserting the expansion at the horizon, Eqs. (67), yields the dimensionless surface gravity

$$\kappa_{\text{sg}} = \frac{f_2(\theta)}{x_H \sqrt{m_2(\theta)}} . \quad (73)$$

As seen from Eq. (68), κ_{sg} is indeed constant on the horizon, as required by the zeroth law of black hole mechanics. The dimensionless temperature T of the black hole is proportional to the surface gravity,

$$T = \frac{\kappa_{\text{sg}}}{2\pi} . \quad (74)$$

E. Electric and magnetic charge

A gauge-invariant definition of the electromagnetic field strength tensor is given by the 't Hooft tensor [16]

$$\mathcal{F}_{\mu\nu} = \text{Tr} \left\{ \hat{\Phi} F_{\mu\nu} - \frac{i}{2e} \hat{\Phi} D_\mu \hat{\Phi} D_\nu \hat{\Phi} \right\} = \hat{\Phi}^a F_{\mu\nu}^a + \frac{1}{e} \epsilon_{abc} \hat{\Phi}^a D_\mu \hat{\Phi}^b D_\nu \hat{\Phi}^c , \quad (75)$$

where $\hat{\Phi}$ is the normalized Higgs field, $|\hat{\Phi}|^2 = (1/2)\text{Tr} \hat{\Phi}^2 = \sum_a (\hat{\Phi}^a)^2 = 1$.

The 't Hooft tensor yields the electric current j_{el}^ν

$$\nabla_\mu \mathcal{F}^{\mu\nu} = -4\pi j_{\text{el}}^\nu , \quad (76)$$

and the magnetic current j_{mag}^ν

$$\nabla_\mu {}^* \mathcal{F}^{\mu\nu} = 4\pi j_{\text{m}}^\nu , \quad (77)$$

where ${}^* \mathcal{F}$ represents the dual field strength tensor.

The electric charge \mathcal{Q} is given by

$$\mathcal{Q} = \frac{Q}{e} = \frac{1}{4\pi} \int_{S_2} {}^* \mathcal{F}_{\theta\varphi} d\theta d\varphi , \quad (78)$$

where the integral is evaluated at spatial infinity.

To define the magnetic charge, we rewrite the 't Hooft tensor as

$$\mathcal{F}_{\mu\nu} = \partial_\mu \mathcal{A}_\nu - \partial_\nu \mathcal{A}_\mu - \frac{i}{2e} \text{Tr} \left\{ \hat{\Phi} \partial_\mu \hat{\Phi} \partial_\nu \hat{\Phi} \right\} , \quad (79)$$

with $\mathcal{A}_\mu = \text{Tr} \left\{ \hat{\Phi} A_\mu \right\}$. Now it follows from Eq. (77) that the magnetic current j_{m}^σ and the topological current k^σ are related by

$$j_{\text{m}}^\sigma = \frac{i}{16\pi e} \epsilon^{\sigma\rho\mu\nu} \text{Tr} \left\{ \partial_\rho \hat{\Phi} \partial_\mu \hat{\Phi} \partial_\nu \hat{\Phi} \right\} = \frac{1}{e} k^\sigma . \quad (80)$$

For globally regular solutions the integration of the magnetic charge density reduces to a surface integral at spacial infinity which yields

$$\mathcal{P} = \frac{n}{e} \varepsilon .$$

For black hole solutions we define the magnetic charge by its value on the horizon plus a volume integral,

$$\mathcal{P} = \mathcal{P}_H + \int_{\Sigma} (-j_{m\mu} n^\mu) dV = \mathcal{P}_H + \int_{r_H}^{\infty} j_m^0 \sqrt{-g} dr d\theta d\varphi , \quad (81)$$

where Σ now denotes an asymptotically flat spacelike hypersurface bounded by the horizon H , dV is the natural volume element on Σ , and n^μ is normal to Σ with $n_\mu n^\mu = -1$.

In order to define the horizon magnetic charge we consider the normalized Higgs field at the horizon as a map between two two-dimensional spheres, which can be characterised by a topological number,

$$N_H = \frac{-i}{16\pi} \int_H \text{Tr} \left\{ \hat{\Phi} d\hat{\Phi} \wedge d\hat{\Phi} \right\} , \quad (82)$$

and obtain $\mathcal{P}_H = N_H/e$. For the evaluation of the magnetic charge we note that the volume integral reduces to a surface integral. Its contribution from the horizon cancels exactly the horizon magnetic charge, and the contribution from the asymptotic region yields $\mathcal{P} = \varepsilon n/e$. Note that for odd m the horizon magnetic charge is either equal to the magnetic charge or to its negative value, depending on how often the Higgs field function Φ_1 changes sign on the symmetry axis. For even m both the magnetic charge and the horizon magnetic charge are zero.

F. Physical Interpretation of ν

The quantity ν is related to the asymptotic behaviour of the gauge potential A_0 , and therefore it is not defined in a gauge-invariant way. To find a physical interpretation of ν we apply a gauge transformation that leads to an asymptotically trivial gauge potential (for even m). Such a gauge transformation is given by

$$U = e^{i\nu t \tau_z/2} e^{im\theta \tau_\varphi^{(n)}/2} .$$

The transformed gauge potential and Higgs field are found to be

$$\begin{aligned} A_\mu dx^\mu = & \left(\left[\bar{B}_1 - \nu - n \frac{\omega}{x} (\cos(m\theta) - 1) \right] \frac{\tau_z}{2e} + \left[B_2 + n \frac{\omega}{x} \sin(m\theta) \right] \frac{\tau_\rho^{(n,\nu t)}}{2e} \right) dt \\ & + A_\varphi (d\varphi - \frac{\omega}{x} dt) + \left(\frac{H_1}{x} dx + (1 - H_2 - m) d\theta \right) \frac{\tau_\varphi^{(n,\nu t)}}{2e} , \end{aligned} \quad (83)$$

with

$$A_\varphi = -n \sin \theta \left(\left[H_3 + \frac{\cos(m\theta) - 1}{\sin \theta} \right] \frac{\tau_z}{2e} + \left[1 - H_4 - \frac{\sin(m\theta)}{\sin \theta} \right] \frac{\tau_\rho^{(n,\nu t)}}{2e} \right) , \quad (84)$$

and

$$\Phi = \left(\Phi_1 \tau_z + \Phi_2 \tau_\rho^{(n,\nu t)} \right) , \quad (85)$$

respectively, where now

$$\begin{aligned} \tau_\rho^{(n,\nu t)} &= \cos(n\varphi - \nu t) \tau_x + \sin(n\varphi - \nu t) \tau_y , \\ \tau_\varphi^{(n,\nu t)} &= -\sin(n\varphi - \nu t) \tau_x + \cos(n\varphi - \nu t) \tau_y . \end{aligned}$$

We observe that in this gauge the fields are explicitly time dependent and rotate in internal space about the τ_z direction. The quantity ν is exactly the rotation frequency.

In the presence of magnetic charge, i.e. odd m , the transformed gauge potential is singular on the negative z axis. However, the physical interpretation of ν does not change.

IV. STATIONARY GLOBALLY REGULAR EYMHD SOLUTIONS

A. Global Charges

Mass, angular momentum and dilaton charge

We begin by recalling the general expressions [24] for the global mass

$$\mathcal{M} = \frac{1}{4\pi G} \int_{\Sigma} R_{\mu\nu} n^{\mu} \xi^{\nu} dV , \quad (86)$$

and the global angular momentum

$$\mathcal{J} = -\frac{1}{8\pi G} \int_{\Sigma} R_{\mu\nu} n^{\mu} \eta^{\nu} dV . \quad (87)$$

Here Σ denotes an asymptotically flat spacelike hypersurface, n^{μ} is normal to Σ with $n_{\mu} n^{\mu} = -1$, and dV is the natural volume element on Σ [24].

Now we express the Ricci tensor in terms of the Yang-Mills, Higgs and dilaton fields, using the Einstein equations, the definition of the stress energy tensor and the Lagrangian

$$\begin{aligned} \frac{1}{8\pi G} R_{\mu\nu} &= \partial_{\mu} \Psi \partial_{\nu} \Psi + 2e^{2\kappa\Psi} \text{Tr}(F_{\mu}^{\alpha} F_{\nu\alpha}) - \frac{1}{2} e^{2\kappa\Psi} \text{Tr}(F_{\rho\sigma} F^{\rho\sigma}) g_{\mu\nu} \\ &\quad + \frac{1}{2} \text{Tr}(D_{\mu} \Phi D_{\nu} \Phi) + \frac{\lambda}{8} e^{-2\kappa\Psi} \text{Tr}(\Phi^2 - v^2)^2 g_{\mu\nu} . \end{aligned} \quad (88)$$

Next we replace the third and the last term in Eq. (88) via the dilaton equation

$$\frac{1}{8\pi G} R_{\mu\nu} = \partial_{\mu} \Psi \partial_{\nu} \Psi + 2e^{2\kappa\Psi} \text{Tr}(F_{\mu}^{\alpha} F_{\nu\alpha}) + \frac{1}{2} \text{Tr}(D_{\mu} \Phi D_{\nu} \Phi) - \frac{1}{2\kappa} \frac{1}{\sqrt{-g}} \partial_{\lambda} (\sqrt{-g} \partial^{\lambda} \Psi) g_{\mu\nu} . \quad (89)$$

Since ξ and η are Killing vector fields and since η is tangential to Σ , we have

$$\xi^{\mu} \partial_{\mu} \Psi = 0 , \quad \eta^{\mu} \partial_{\mu} \Psi = 0 , \quad n^{\mu} \eta^{\nu} g_{\mu\nu} = 0 , \quad (90)$$

and consequently,

$$\frac{1}{8\pi G} R_{\mu\nu} n^{\mu} \xi^{\nu} = 2e^{2\kappa\Psi} \text{Tr}(F_{\mu}^{\alpha} F_{\nu\alpha}) n^{\mu} \xi^{\nu} + \frac{1}{2} \text{Tr}(D_{\mu} \Phi D_{\nu} \Phi) n^{\mu} \xi^{\nu} - \frac{1}{2\kappa} \frac{1}{\sqrt{-g}} \partial_{\lambda} (\sqrt{-g} \partial^{\lambda} \Psi) n^{\mu} \xi_{\mu} , \quad (91)$$

$$\frac{1}{8\pi G} R_{\mu\nu} n^{\mu} \eta^{\nu} = 2e^{2\kappa\Psi} \text{Tr}(F_{\mu}^{\alpha} F_{\nu\alpha}) n^{\mu} \eta^{\nu} + \frac{1}{2} \text{Tr}(D_{\mu} \Phi D_{\nu} \Phi) n^{\mu} \eta^{\nu} . \quad (92)$$

We now define the dilaton charge \mathcal{D} via

$$\int_{\Sigma} \frac{1}{\sqrt{-g}} \partial_{\lambda} (\sqrt{-g} \partial^{\lambda} \Psi) n^{\mu} \xi_{\mu} dV = -4\pi \mathcal{D} . \quad (93)$$

Making use of the dilaton charge \mathcal{D} , we obtain for the mass \mathcal{M}

$$\mathcal{M} = 4 \int_{\Sigma} \{ e^{2\kappa\Psi} \text{Tr}(F_{\mu}^{\alpha} F_{\nu\alpha}) n^{\mu} \xi^{\nu} \} dV + \int_{\Sigma} \{ \text{Tr}(D_{\mu} \Phi D_{\nu} \Phi) n^{\mu} \xi^{\nu} \} dV + \frac{4\pi}{\kappa} \mathcal{D} , \quad (94)$$

while the angular momentum \mathcal{J} is given by

$$\mathcal{J} = -2 \int_{\Sigma} \{ e^{2\kappa\Psi} \text{Tr}(F_{\mu}^{\alpha} F_{\nu\alpha}) n^{\mu} \eta^{\nu} \} dV - \frac{1}{2} \int_{\Sigma} \{ \text{Tr}(D_{\mu} \Phi D_{\nu} \Phi) n^{\mu} \eta^{\nu} \} dV . \quad (95)$$

To evaluate the integrals in Eq. (94) and Eq. (95) we use local coordinates (t, r, θ, φ) . In these coordinates

$$n^{\mu} = -\sqrt{f} g^{0\mu} , \quad \xi^{\mu} = (1, 0, 0, 0) , \quad \eta^{\mu} = (0, 0, 0, 1) , \quad dV = \frac{1}{\sqrt{f}} \sqrt{-g} dr d\theta d\varphi , \quad (96)$$

and we obtain

$$\mathcal{M} - \frac{4\pi}{\kappa} \mathcal{D} = \mathcal{I}_{\mathcal{M}} \equiv -4 \int_{\Sigma} e^{2\kappa\Psi} \text{Tr} [F_{0\mu} F^{0\mu}] \sqrt{-g} dr d\theta d\varphi - \int_{\Sigma} \text{Tr} [D_0 \Phi D^0 \Phi] \sqrt{-g} dr d\theta d\varphi , \quad (97)$$

$$\mathcal{J} = \mathcal{I}_{\mathcal{J}} \equiv 2 \int_{\Sigma} e^{2\kappa\Psi} \text{Tr} [F_{\varphi\mu} F^{0\mu}] \sqrt{-g} dr d\theta d\varphi + \frac{1}{2} \int_{\Sigma} \text{Tr} [D_{\varphi} \Phi D^0 \Phi] \sqrt{-g} dr d\theta d\varphi , \quad (98)$$

defining the integrals $\mathcal{I}_{\mathcal{M}}$ and $\mathcal{I}_{\mathcal{J}}$.

To evaluate the integrals $\mathcal{I}_{\mathcal{M}}$, Eq. (97) and $\mathcal{I}_{\mathcal{J}}$, Eq. (98), we make use of the symmetry relations, Eqs. (15) [20],

$$F_{\mu 0} = \hat{D}_{\mu} A_0 , \quad F_{\mu\varphi} = \hat{D}_{\mu} (A_{\varphi} - W_{\eta}) , \quad (99)$$

where $\hat{D}_{\mu} \equiv \partial_{\mu} + ie[A_{\mu}, \cdot]$. The integrals then read

$$\mathcal{I}_{\mathcal{M}} = +4 \int_{\Sigma} e^{2\kappa\Psi} \text{Tr} [\hat{D}_{\mu} A_0 F^{0\mu}] \sqrt{-g} dr d\theta d\varphi - \int_{\Sigma} \text{Tr} [D_0 \Phi D^0 \Phi] \sqrt{-g} dr d\theta d\varphi , \quad (100)$$

$$\mathcal{I}_{\mathcal{J}} = -2 \int_{\Sigma} e^{2\kappa\Psi} \text{Tr} [\hat{D}_{\mu} (A_{\varphi} - W_{\eta}) F^{0\mu}] \sqrt{-g} dr d\theta d\varphi + \frac{1}{2} \int_{\Sigma} \text{Tr} [D_{\varphi} \Phi D^0 \Phi] \sqrt{-g} dr d\theta d\varphi . \quad (101)$$

Adding zero to the above integrals, in the form of the gauge field equation of motion for the zero component, we obtain

$$\mathcal{I}_{\mathcal{M}} = 4 \int_{\Sigma} \text{Tr} [\hat{D}_{\mu} \{A_0 e^{2\kappa\Psi} F^{0\mu} \sqrt{-g}\}] dr d\theta d\varphi + \int_{\Sigma} \text{Tr} [ie A_0 [\Phi, D^0 \Phi] - D_0 \Phi D^0 \Phi] \sqrt{-g} dr d\theta d\varphi , \quad (102)$$

$$\begin{aligned} \mathcal{I}_{\mathcal{J}} &= -2 \int_{\Sigma} \text{Tr} [\hat{D}_{\mu} \{(A_{\varphi} - W_{\eta}) e^{2\kappa\Psi} F^{0\mu} \sqrt{-g}\}] dr d\theta d\varphi \\ &\quad - \frac{1}{2} \int_{\Sigma} \text{Tr} [ie (A_{\varphi} - W_{\eta}) [\Phi, D^0 \Phi] - D_{\varphi} \Phi D^0 \Phi] \sqrt{-g} dr d\theta d\varphi . \end{aligned} \quad (103)$$

Making use of the explicit form of the ansatz, exploiting in particular Eq. (14), we see, that for both \mathcal{I}_M and \mathcal{I}_J the second integral vanishes identically, leaving only the first integral to be analyzed further.

Since the trace of a commutator vanishes, we now replace the derivative \hat{D}_{μ} by the partial derivative ∂_{μ} in the remaining integrals,

$$I_M = 4 \int_{\Sigma} \text{Tr} [\partial_{\mu} \{A_0 e^{2\kappa\Psi} F^{0\mu} \sqrt{-g}\}] dr d\theta d\varphi , \quad (104)$$

$$\mathcal{I}_J = -2 \int_{\Sigma} \text{Tr} [\partial_{\mu} \{(A_{\varphi} - W_{\eta}) e^{2\kappa\Psi} F^{0\mu} \sqrt{-g}\}] dr d\theta d\varphi , \quad (105)$$

and employ the divergence theorem. The θ -term vanishes, since $\sqrt{-g}$ vanishes at $\theta = 0$ and $\theta = \pi$, and the φ -term vanishes, since the integrands at $\varphi = 0$ and $\varphi = 2\pi$ coincide, thus we are left with

$$\mathcal{I}_M = 4 \int \text{Tr} [A_0 e^{2\kappa\Psi} F^{0r} \sqrt{-g}]|_0^{\infty} d\theta d\varphi , \quad (106)$$

$$\mathcal{I}_J = -2 \int \text{Tr} [(A_{\varphi} - W_{\eta}) e^{2\kappa\Psi} F^{0r} \sqrt{-g}]|_0^{\infty} d\theta d\varphi . \quad (107)$$

Since the integrands vanish at the origin, the only contributions to \mathcal{I}_M and \mathcal{I}_J come from infinity. At infinity the asymptotic expansion yields to lowest order

$$\begin{aligned} F^{0r} \sqrt{-g} &= Q \sin \theta \frac{\tau_r^{(n,m)}}{2e} + o(1) , \\ A_0 &= \tilde{\nu} \frac{\tau_r^{(n,m)}}{2e} + o(1) , \\ A_{\varphi} &= -n \sin \theta \left[\frac{\cos(\varepsilon\theta) - \cos(m\theta)}{\sin \theta} \frac{\tau_r^{(n,m)}}{2e} + \frac{\sin(m\theta)}{\sin \theta} \frac{\tau_{\theta}^{(n,m)}}{2e} \right] + o(1) , \end{aligned} \quad (108)$$

where

$$\tilde{\nu} = \frac{e\alpha}{\sqrt{4\pi G}}\nu, \quad \varepsilon = \frac{1}{2}(1 - (-1)^m). \quad (109)$$

The integrals \mathcal{I}_M and \mathcal{I}_J are then given by

$$\mathcal{I}_M = \frac{8\pi\tilde{\nu}Q}{e^2}, \quad \mathcal{I}_J = \frac{4\pi nQ}{e^2}(1 - \varepsilon), \quad (110)$$

yielding for the mass \mathcal{M} and the angular momentum \mathcal{J}

$$\mathcal{M} = \frac{4\pi}{\kappa}\mathcal{D} + \frac{8\pi\tilde{\nu}Q}{e^2}, \quad \mathcal{J} = \frac{4\pi nQ}{e^2}(1 - \varepsilon). \quad (111)$$

Returning to dimensionless variables, and noting that

$$\mathcal{M} = \frac{\sqrt{4\pi G}}{e\alpha G}\mu, \quad \mathcal{J} = \frac{4\pi}{e^2\alpha^2}\zeta, \quad \mathcal{D} = \frac{D}{e}, \quad (112)$$

we obtain the mass formula

$$\mu = \alpha^2 \frac{D}{\gamma} + 2\alpha^2 \nu Q \iff M = \frac{D}{\gamma} + 2\nu Q, \quad (113)$$

and the quantization condition for the angular momentum Eq. (1)

$$\zeta = \alpha^2 nQ(1 - \varepsilon) \iff J = nQ(1 - \varepsilon).$$

B. Effective Action

To address the dependence of the globally regular solutions on the coupling constant α , we now consider the effective action \mathcal{S}^{eff} . In particular, we explain the qualitatively different dependence of the mass M for static and for stationary solutions. This concerns only such types of regular solutions where two branches of solutions exist.

For static solutions, the mass M exhibits a “spike” at the maximal value of the coupling α_{max} , where the branches merge and end [9, 11, 17]. The tangent of the mass w.r.t. α must be the same for both branches at α_{max} [23]. In contrast, for stationary solutions, the mass M exhibits a “loop” in the vicinity of the maximal value of the coupling α_{max} [22]. Here the tangent of the mass w.r.t. α diverges at α_{max} . The loop is associated with a critical value of α , where the two mass branches cross.

Effective action and mass

Let us begin by defining the effective action \mathcal{S}^{eff} ,

$$\mathcal{S}^{\text{eff}} = \int \left(\frac{\hat{R}}{16\pi G} + L_M \right) \sqrt{-g} d^3x, \quad (114)$$

with the gravitational effective Lagrangian

$$\frac{\hat{R}}{16\pi G} = \frac{1}{16\pi G} \left(R - \frac{\partial_\mu \Delta^\mu}{\sqrt{-g}} \right) \quad (115)$$

and the matter Lagrangian L_M Eq. (5)

$$L_M = -\frac{1}{2}\partial_\mu \Psi \partial^\mu \Psi - \frac{1}{2}e^{2\kappa\Psi} \text{Tr}(F_{\mu\nu} F^{\mu\nu}) - \frac{1}{4}\text{Tr}(D_\mu \Phi D^\mu \Phi) - \frac{\lambda}{8}e^{-2\kappa\Psi} \text{Tr}(\Phi^2 - v^2)^2.$$

The divergence term $\partial_\mu \Delta^\mu$ in the gravitational effective Lagrangian ensures that the variational principle of the effective action \mathcal{S}^{eff} w.r.t. the functions of the ansatz yields the proper set of field equations. For our particular ansatz of the metric Δ^μ is given by

$$\Delta^\mu = \sqrt{l} \sin \theta \ (0, r^2 \partial_r, \partial_\theta, 0) \ln \frac{f}{ml}. \quad (116)$$

Reexpressing the curvature scalar R via the Einstein equations,

$$\frac{R}{8\pi G} = \partial_\mu \Psi \partial^\mu \Psi + \frac{1}{2} \text{Tr} (D_\mu \Phi D^\mu \Phi) + \frac{\lambda}{2} e^{-2\kappa \Psi} \text{Tr} (\Phi^2 - v^2)^2, \quad (117)$$

then leads to the effective action

$$\mathcal{S}^{\text{eff}} = \int \left(-\frac{1}{16\pi G} \frac{\partial_\mu \Delta^\mu}{\sqrt{-g}} - \frac{1}{2} e^{2\kappa \Psi} \text{Tr} (F_{\mu\nu} F^{\mu\nu}) + \frac{\lambda}{8} e^{-2\kappa \Psi} \text{Tr} (\Phi^2 - v^2)^2 \right) \sqrt{-g} d^3 x. \quad (118)$$

Analogously to the derivation of the mass formula, we next replace the two matter terms in \mathcal{S}^{eff} via the equation of motion of the dilaton field and obtain

$$\mathcal{S}^{\text{eff}} = -\frac{1}{16\pi G} \int \partial_\mu \Delta^\mu dr d\theta d\varphi + \frac{1}{2\kappa} \int \frac{1}{\sqrt{-g}} \partial_\lambda (\sqrt{-g} \partial^\lambda \Psi) \sqrt{-g} dr d\theta d\varphi. \quad (119)$$

Since in the local coordinates the second integral agrees with the integral for the dilaton charge Eq. (93) and since the θ -term in the first integral vanishes, the effective action becomes

$$\mathcal{S}^{\text{eff}} = -\frac{1}{16\pi G} \int \Delta^r d\theta d\varphi - \frac{4\pi \mathcal{D}}{2\kappa}. \quad (120)$$

The remaining integral is evaluated with help of the asymptotic expansion of the metric functions, leading to

$$\mathcal{S}^{\text{eff}} = -\frac{1}{2} \left(\mathcal{M} + \frac{4\pi \mathcal{D}}{\kappa} \right), \quad (121)$$

which can be rewritten via the mass formula for the regular solutions Eq. (111)

$$\mathcal{S}^{\text{eff}} = -\mathcal{M} + \frac{4\pi \tilde{\nu} Q}{e^2}. \quad (122)$$

Defining finally the dimensionless effective action S^{eff}

$$S^{\text{eff}} = \frac{\sqrt{4\pi G}}{e\alpha G} \alpha^2 \mathcal{S}^{\text{eff}}, \quad (123)$$

we obtain

$$S^{\text{eff}} = -(M - \nu Q) \quad (124)$$

or equivalently

$$S^{\text{eff}} = -\left(\frac{D}{\gamma} + \nu Q \right). \quad (125)$$

Note, that Eq. (124) for the effective action remains true when the dilaton decouples, i.e. Eq. (124) also holds for EYMH solutions.

Dependence of the effective action S^{eff} and the mass M on the coupling constant α

To address the dependence of the effective action S^{eff} on the coupling constant α we first make the α -dependence more explicit. To this end, we express all quantities in the effective action in dimensionless quantities. Denoting the by the divergence term corrected dimensionless curvature scalar \bar{R} (i.e. $\hat{R} \rightarrow \bar{R}$), the dimensionless matter Lagrangian \bar{L}_M , and the dimensionless determinant of the metric $-\bar{g}$, and we obtain the dimensionless effective action,

$$S^{\text{eff}} = \frac{1}{4\pi} \left[\frac{1}{4\alpha^2} \int \bar{R} \sqrt{-\bar{g}} d^3 x + \int \bar{L}_M \sqrt{-\bar{g}} d^3 x \right]. \quad (126)$$

We now take the derivative of S^{eff} w.r.t. α , taking into account that the metric and matter functions, abbreviated by X_i , implicitly also depend on α . The derivative has thus two terms

$$\frac{dS^{\text{eff}}}{d\alpha} = \frac{1}{4\pi} \left[-\frac{1}{2\alpha^3} \int \bar{R} \sqrt{-\bar{g}} d^3 x \right] + \frac{1}{16\pi\alpha^2} \int \underbrace{\left\{ \frac{\partial(\bar{R}\sqrt{-\bar{g}})}{\partial X_i} + 4\alpha^2 \frac{\partial(\bar{L}_M\sqrt{-\bar{g}})}{\partial X_i} \right\}}_{=0 \text{ equations of motion}} \frac{\partial X_i}{\partial \alpha} d^3 x, \quad (127)$$

where the second term vanishes for solutions of the equations of motion, and we are left with

$$\frac{dS^{\text{eff}}}{d\alpha} = -\frac{1}{8\pi\alpha^3} \int \bar{R}\sqrt{-g}d^3x . \quad (128)$$

From Eq. (128) we conclude, that the effective action S^{eff} must exhibit a “spike” at the maximal value of the coupling α_{max} , where the two branches of solutions merge and end, since the tangent w.r.t. α must be the same for both branches at α_{max} . This conclusion holds for stationary solutions, as well as for static solutions.

Let us now address the mass M , which is related to the effective action via Eq. (124), i.e.

$$M = -S^{\text{eff}} + \nu Q . \quad (129)$$

Considering the derivative of the mass M w.r.t. α , keeping ν fixed, we obtain

$$\frac{dM}{d\alpha} = -\frac{dS^{\text{eff}}}{d\alpha} + \nu \frac{dQ}{d\alpha} . \quad (130)$$

Clearly, the crucial difference between the tangent of the mass and the tangent of the effective action resides in the last term, containing the derivative of the electric charge Q w.r.t. α . It is this term which allows different tangents on both branches at α_{max} . In fact, for stationary (non-static) solutions, we observe that this last term diverges at α_{max} , yielding a divergent tangent also for the mass, as required for a “loop” associated with both mass branches in the vicinity of α_{max} . For static solutions, on the other hand, the mass M always exhibits a “spike” at the maximal value of the coupling, α_{max} , since the electric charge Q vanishes.

Note however, that we assumed so far that the value of ν is fixed by a boundary condition, $\bar{B}_1|_{\infty} = \nu$. This is in contrast to the case where the electric charge is kept fixed by the boundary condition at infinity,

$$x^2 \partial_x \bar{B}_1|_{\infty} = Q . \quad (131)$$

In this case the quantity ν is allowed to vary and its value is adjusted by the numerical procedure. More formally, in the variation of the effective action a boundary term shows up, which evaluates to $Q\delta\nu$. Therefore the field equations are obtained from the modified effective action,

$$\delta (S^{\text{eff}} - Q\nu)|_Q = 0 \quad (132)$$

As a consequence, if α is varied for fixed electric charge Q ,

$$\frac{d}{d\alpha} (S^{\text{eff}} - Q\nu)|_Q = -\frac{1}{8\pi\alpha^3} \int \bar{R}\sqrt{-g}d^3x , \quad (133)$$

since the variation of the modified effective action with respect to the fields vanishes. On the other hand, since $S^{\text{eff}} - Q\nu = -M$, we find

$$\frac{dM}{d\alpha} = \frac{1}{8\pi\alpha^3} \int \bar{R}\sqrt{-g}d^3x . \quad (134)$$

Thus for fixed electric charge it is the mass that exhibits a spike.

C. Numerical results

We solve the set of thirteen coupled non-linear elliptic partial differential equations numerically [33], subject to the above boundary conditions, requiring the solutions to be regular at the origin. We employ compactified dimensionless coordinates, $\bar{x} = x/(1+x)$. The numerical calculations, based on the Newton-Raphson method, are performed with help of the program FIDISOL [33]. The equations are discretized on a non-equidistant grid in \bar{x} and θ . Typical grids used have sizes 100×20 , covering the integration region $0 \leq \bar{x} \leq 1$ and $0 \leq \theta \leq \pi/2$. (See [12, 30] and [33] for further details on the numerical procedure.)

For given coupling constants α , β and γ , the stationary globally regular solutions then depend on the parameter ν , specifying the time component of the gauge potential at infinity, and on the integers m and n . (In principle, the solutions can further depend on the node number of the gauge potential functions k , labelling the radial excitations. However, we here focus on the lowest mass solutions.)

We now illustrate the above relations for the mass and the effective action with numerical results presented in Figs. 1 and 2. We first consider stationary gravitating dyon solutions with $m = 1$. Dyons with magnetic charge $n = 1$ are spherically symmetric [31], dyons with higher magnetic charge are axially symmetric [22].

In Fig. 1a we exhibit the mass M of dyons with magnetic charge $n = 1, 2$ and 3 at a fixed value of ν versus the coupling constant α (at $\beta = \gamma = 0$). In each case, a first branch of gravitating dyons emerges from the corresponding flat space solution at $\alpha = 0$ and extends up to a maximal value of the coupling constant, α_{\max} , beyond which no dyon solutions exist. For the $n = 1$ dyons we observe a second branch of solutions in the vicinity of the maximal value of α . This second branch ends at a critical value of α , where the branch of non-Abelian solutions merges with the corresponding branch of extremal Reissner-Nordström solutions [9, 31]. For $n > 1$ dyon solutions, numerical accuracy does not allow us to discern the existence of two branches.

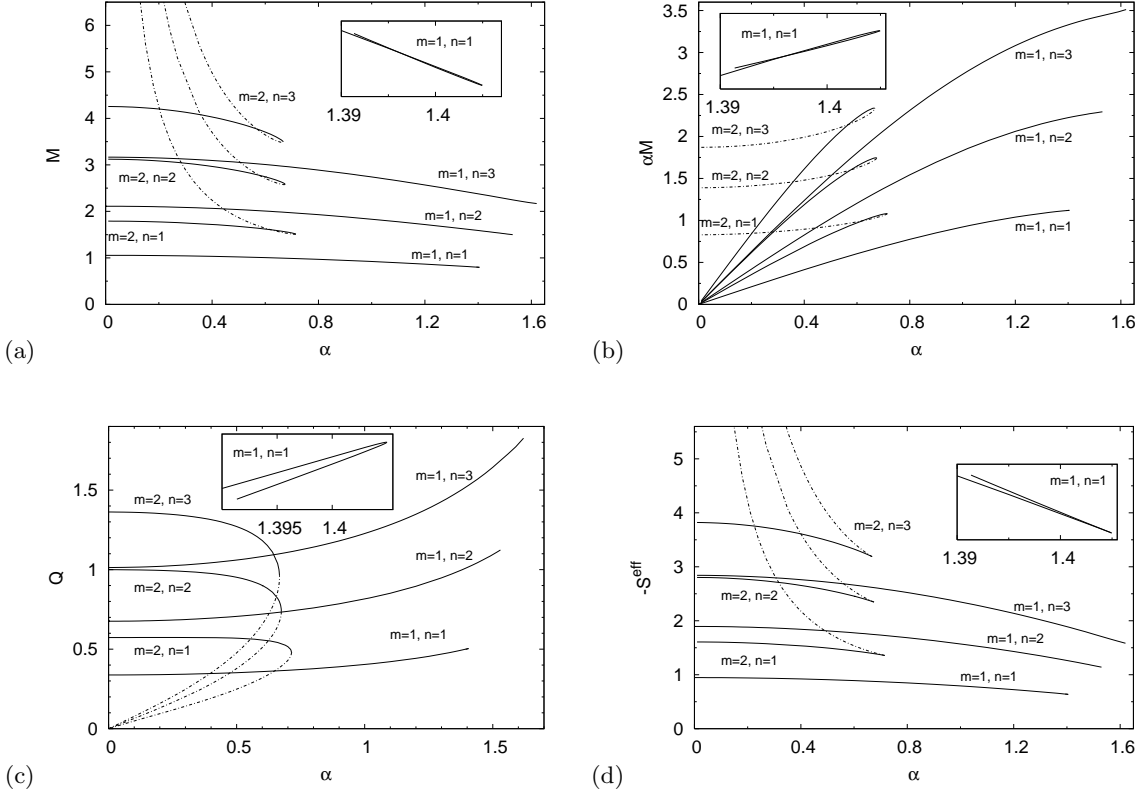


Fig. 1 The mass M (a), the scaled mass αM (b), the electric charge Q (c), and the effective action S^{eff} (d) are shown versus the coupling constant α for dyon solutions with $m = 1$, $n = 1, 2, 3$ and for electrically charged monopole-antimonopole resp. vortex ring solutions with $m = 2$, $n = 1, 2, 3$ at $\nu = 0.32$, $\beta = \gamma = 0$ (first branch: solid, second branch: dot-dashed).

Besides dyons, we also exhibit in Fig. 1a the mass of electrically charged monopole-antimonopole pair resp. vortex ring solutions, which have $m = 2$ and $n = 1, 2$ and 3 . For these solutions always two branches of solutions exist. Again, the first branch emerges from the respective flat space solution, and extends up to a maximal value of α , where it merges with the second branch. But the second branch now extends back to $\alpha \rightarrow 0$. The mass diverges on the second branch in the limit $\alpha \rightarrow 0$. But considering the scaled mass αM instead, exhibited in Fig. 1b, one realizes, that in the limit $\alpha \rightarrow 0$ a globally regular EYM solution [15, 30] is reached (after rescaling) [17]. Clearly, the electric charge Q , exhibited in Fig. 1c, also tends to zero on the second branch in the limit $\alpha \rightarrow 0$, yielding non-rotating limiting EYM solutions, in agreement with previous results on globally regular EYM solutions [19, 20].

We exhibit the effective action S^{eff} for the same set of solutions in Fig. 1d. As predicted above, the effective action exhibits a “spike” close to the maximal value of the coupling constant α , whenever two branches of solutions are present. The mass, in contrast, exhibits a “loop” for these stationary non-static solutions close to α_{\max} . At α_{\max} the tangent of the mass diverges, since the tangent of the electric charge diverges there. If we consider branches of static solutions instead, the mass exhibits a “spike” close to α_{\max} [9, 17].

While in the solutions of Fig. 1 the dilaton is decoupled since $\gamma = 0$, we consider in Fig. 2 the dependence of the solutions on the dilaton coupling constant γ . In Fig. 2a we exhibit the mass and the effective action of electrically

charged monopole-antimonopole pair solutions ($m = 2, n = 1, \nu = 0.32, \beta = 0$) versus the coupling constant α at fixed dilaton coupling constant $\gamma = 0.4$. As predicted, we observe a “spike” for the effective action and a “loop” for the mass. Note, that the mass of these EYMHD solutions does not diverge on the second branch in the limit $\alpha \rightarrow 0$, since a YMHD solution is approached. We note, that the effect of a dilaton on globally regular solutions is very similar to the effect of gravity [26].

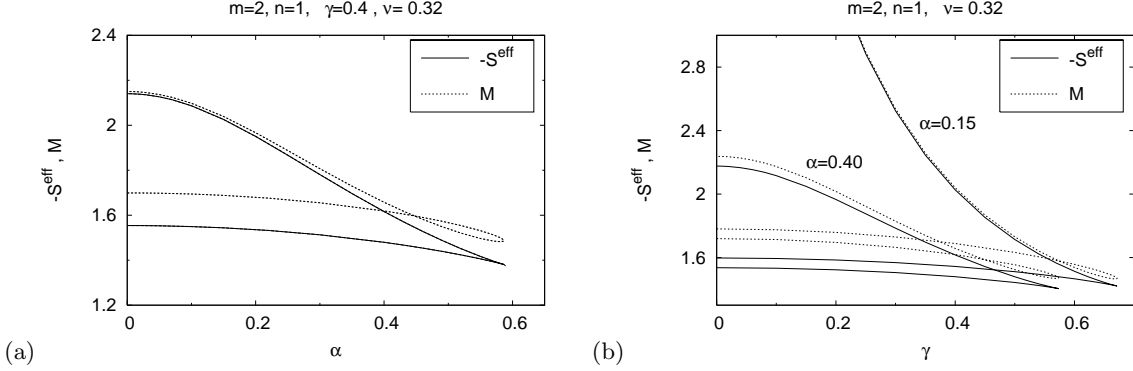


Fig. 2 The mass M and the effective action S^{eff} of electrically charged monopole-antimonopole solutions with $m = 2, n = 1, \nu = 0.32, \beta = 0$ are shown versus the coupling constant α for dilaton coupling constant $\gamma = 0.4$ (a), and versus the dilaton coupling constant γ for coupling constant $\alpha = 0.15$ and 0.40 (b).

The γ -dependence of the solutions (at fixed α) is illustrated in Fig. 2b for electrically charged monopole-antimonopole pair solutions with $m = 2, n = 1, \nu = 0.32$ ($\beta = 0$) for $\alpha = 0.15$ and 0.40 . Again we note, that the effective action exhibits a “spike” while the mass exhibits a “loop”. That this must be the case, can be shown by an argument analogous to the one employed above for the α -dependence. Note, that the mass of these EYMHD solutions does not diverge on the second branch in the limit $\gamma \rightarrow 0$, since EYM solutions are approached in this limit.

We illustrate the globally regular solutions with an example in Fig. 3. We here exhibit the energy density of the matter fields ϵ

$$\epsilon = -\frac{2}{e^2 v^4} \left(T_0^0 - \frac{1}{2} T_\mu^\mu \right) \quad (135)$$

for a monopole-antimonopole pair solution ($m = 2, n = 1$) carrying electric charge and angular momentum ($\nu = 0.32, \alpha = 0.3, \beta = \gamma = 0$). The maxima of the energy density are associated with the location of the magnetic poles on the symmetry axis.

V. ROTATING EYMHD BLACK HOLES

A. Non-Abelian Mass Formula

We now derive the mass formula for stationary axially symmetric EYMHD black hole solutions

$$M = 2\hat{T}\hat{S} + 2\Omega J + \frac{D}{\gamma} + 2\nu Q + 2\tilde{\psi}_{\text{el}} Q(1 - \varepsilon), \quad (136)$$

where \hat{T} and \hat{S} are conveniently scaled dimensionless temperature and entropy, respectively. The derivation is analogous to the derivation of the expressions for the mass and the angular momentum of the globally regular solutions.

Again we begin by recalling the general expressions [24] for the global mass

$$\mathcal{M} = \mathcal{M}_{\text{H}} + \frac{1}{4\pi G} \int_{\Sigma} R_{\mu\nu} n^\mu \xi^\nu dV, \quad (137)$$

and the global angular momentum

$$\mathcal{J} = \mathcal{J}_{\text{H}} - \frac{1}{8\pi G} \int_{\Sigma} R_{\mu\nu} n^\mu \eta^\nu dV, \quad (138)$$

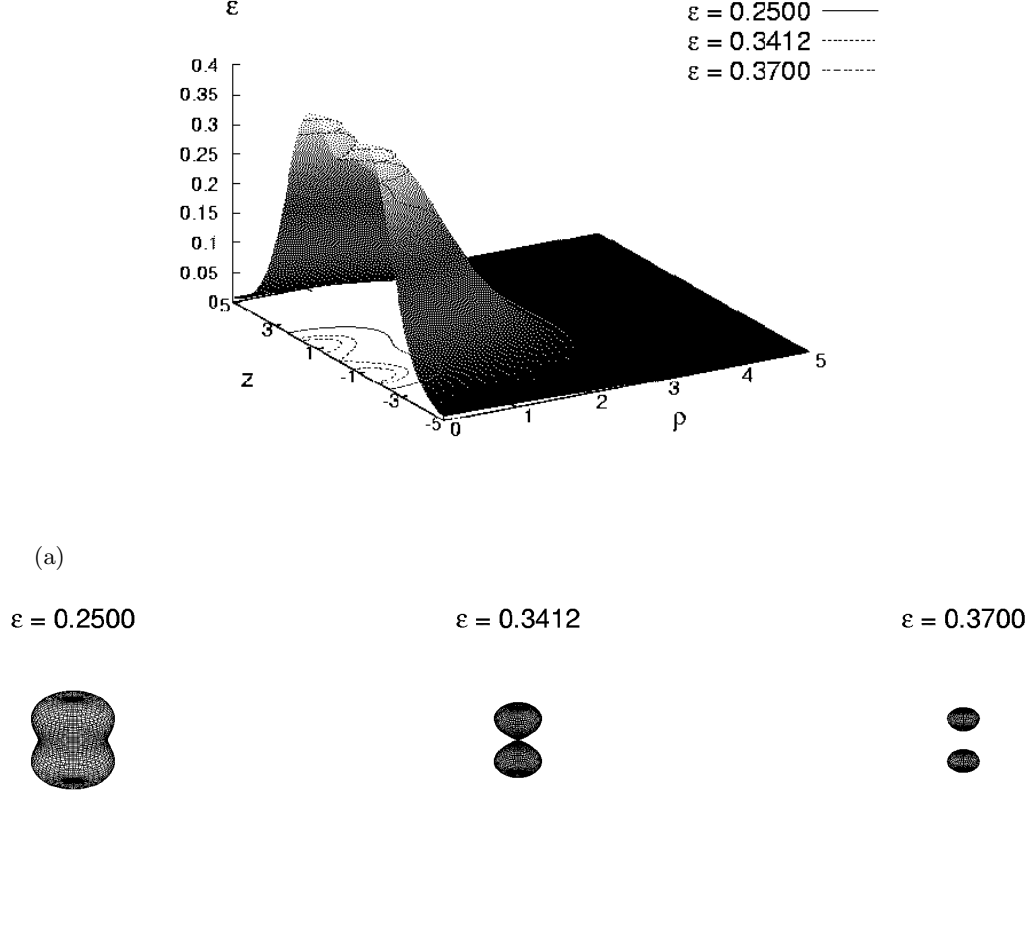


Fig. 3 (a) The energy density of the matter fields for an electrically charged monopole-antimonopole pair solution ($m = 2$, $n = 1$, $\nu = 0.32$, $\alpha = 0.3$, $\beta = \gamma = 0$). (b) Also shown are surfaces of constant energy density.

where Σ now denotes an asymptotically flat spacelike hypersurface bounded by the horizon H , and the horizon mass \mathcal{M}_H [24] and the horizon angular momentum \mathcal{J}_H are given by

$$\mathcal{M}_H = -\frac{1}{8\pi G} \int_H \frac{1}{2} \varepsilon_{\mu\nu\rho\sigma} \nabla^\rho \xi^\sigma dx^\mu dx^\nu = 2TS + 2\frac{\omega_H}{r_H} \mathcal{J}_H \quad (139)$$

$$\mathcal{J}_H = \frac{1}{16\pi G} \int_H \frac{1}{2} \varepsilon_{\mu\nu\rho\sigma} \nabla^\rho \eta^\sigma dx^\mu dx^\nu . \quad (140)$$

Substituting the horizon mass \mathcal{M}_H in Eq. (137) and eliminating the horizon angular momentum \mathcal{J}_H yields for the global mass \mathcal{M}

$$\mathcal{M} = 2TS + 2\frac{\omega_H}{r_H} \mathcal{J} + 2 \left[\frac{1}{8\pi G} \int_\Sigma R_{\mu\nu} n^\mu \xi^\nu dV + \frac{\omega_H}{r_H} \frac{1}{8\pi G} \int_\Sigma R_{\mu\nu} n^\mu \eta^\nu dV \right] . \quad (141)$$

Following now the same steps as for the globally regular solutions, we obtain analogously to Eq. (97) and Eq. (98)

$$\begin{aligned} \mathcal{M} - 2TS - 2\frac{\omega_H}{r_H} \mathcal{J} - \frac{4\pi}{\kappa} \mathcal{D} = \mathcal{I} \equiv & -4 \int_\Sigma e^{2\kappa\Psi} \text{Tr} \left[\left(F_{0\mu} + \frac{\omega_H}{r_H} F_{\varphi\mu} \right) F^{0\mu} \right] \sqrt{-g} dr d\theta d\varphi \\ & - \int_\Sigma \text{Tr} \left[\left(D_0 \Phi + \frac{\omega_H}{r_H} D_\varphi \Phi \right) D^0 \Phi \right] \sqrt{-g} dr d\theta d\varphi , \end{aligned} \quad (142)$$

defining the integral \mathcal{I} .

We next note, that formally, we can express the integral \mathcal{I} in terms of the integrals \mathcal{I}_M Eq. (97) and \mathcal{I}_J Eq. (98), keeping in mind, that Σ is here bounded by the horizon. Thus

$$\mathcal{M} - 2\mathcal{T}S - 2\frac{\omega_H}{r_H}\mathcal{J} - \frac{4\pi}{\kappa}\mathcal{D} = \mathcal{I}_M - 2\frac{\omega_H}{r_H}\mathcal{I}_J . \quad (143)$$

To evaluate \mathcal{I} , we proceed again analogously to the globally regular case, making use of the relations obtained for \mathcal{I}_M and \mathcal{I}_J . Analogously to Eq. (106) and Eq. (107) we are then left with

$$\mathcal{I} = 4 \int \text{Tr} \left[\left(A_0 + \frac{\omega_H}{r_H} (A_\varphi - W_\eta) \right) e^{2\kappa\Psi} F^{0r} \sqrt{-g} \right] \Big|_{r_H}^{\infty} d\theta d\varphi , \quad (144)$$

where the integrand must be evaluated at the horizon and at infinity. Since the electrostatic potential is constant at the horizon (see Eqs. (49)) and

$$\left(A_0 + \frac{\omega_H}{r_H} A_\varphi \right) \Big|_H = \frac{\omega_H}{r_H} W_\eta = -\tilde{\Psi}_{\text{el}} \frac{\tau_z}{2} , \quad (145)$$

the integrand vanishes at the horizon, and the only contribution to \mathcal{I} comes from infinity.

At infinity both expressions have been evaluated in Eq. (110). Thus the mass formula becomes

$$\mathcal{M} - 2\mathcal{T}S - 2\frac{\omega_H}{r_H}\mathcal{J} - \frac{4\pi}{\kappa}\mathcal{D} = \frac{8\pi\tilde{\nu}Q}{e^2} + \frac{8\pi\tilde{\Psi}_{\text{el}}Q}{e}(1 - \varepsilon) . \quad (146)$$

Returning again to dimensionless variables, recalling Eq. (50) and Eq. (112), and noting that

$$\mathcal{T}S = \frac{\sqrt{4\pi G}}{e\alpha G} TS , \quad \frac{\omega_H}{r_H} = \frac{e\alpha}{\sqrt{4\pi G}} \Omega , \quad (147)$$

we obtain the mass formula Eq. (136)

$$\mu = 2TS + 2\Omega\zeta + \alpha^2 \frac{D}{\gamma} + 2\alpha^2\nu Q + 2\alpha^2\tilde{\psi}_{\text{el}}Q(1 - \varepsilon) ,$$

or equivalently

$$M = 2\hat{T}\hat{S} + 2\Omega J + \frac{D}{\gamma} + 2\nu Q + 2\tilde{\psi}_{\text{el}}Q(1 - \varepsilon) ,$$

with scaled dimensionless temperature and entropy, $\hat{T}\hat{S} = TS/\alpha^2$.

This mass formula differs from the EMD and EYMD mass formula Eq. (3) in two respects. First, the last term is present only for magnetically neutral black holes. Second, the fourth term is an additional term, not present for EMD and EYMD black holes. It appears for all electrically charged EYMHD black holes, and has the gauge potential parameter ν entering together with the electric charge. We note, that the first two terms and the last term do not appear in the mass formula Eq. (113) for globally regular solutions. Indeed, when the black hole horizon size is taken to zero, the first term vanishes, and the second and the last term cancel, leaving the mass formula Eq. (113) for globally regular solutions.

B. Numerical Results

The numerical black hole calculations are performed analogously to the calculations of globally regular solutions [33], except that for black hole solutions we employ the compactified dimensionless coordinate $\bar{x} = 1 - (x_H/x)$, and we impose boundary conditions at the regular horizon.

For given coupling constants α , β and γ , the rotating non-Abelian black hole solutions then depend on the horizon radius x_H , and on the rotational velocity of the horizon Ω in addition to the gauge potential parameter ν and the integers m and n .

We exhibit in Fig. 4 an example of a dyonic rotating black hole, which has $m = 1$, $n = 1$, horizon radius $x_H = 0.1$, horizon angular velocity $\Omega = 0.5$, gauge potential parameter $\nu = 0.04$ and the coupling constants $\alpha = 0.3$, $\beta = 0.1$, $\gamma = 0.1$. We again exhibit the energy density of the matter fields ϵ , Eq. (135).

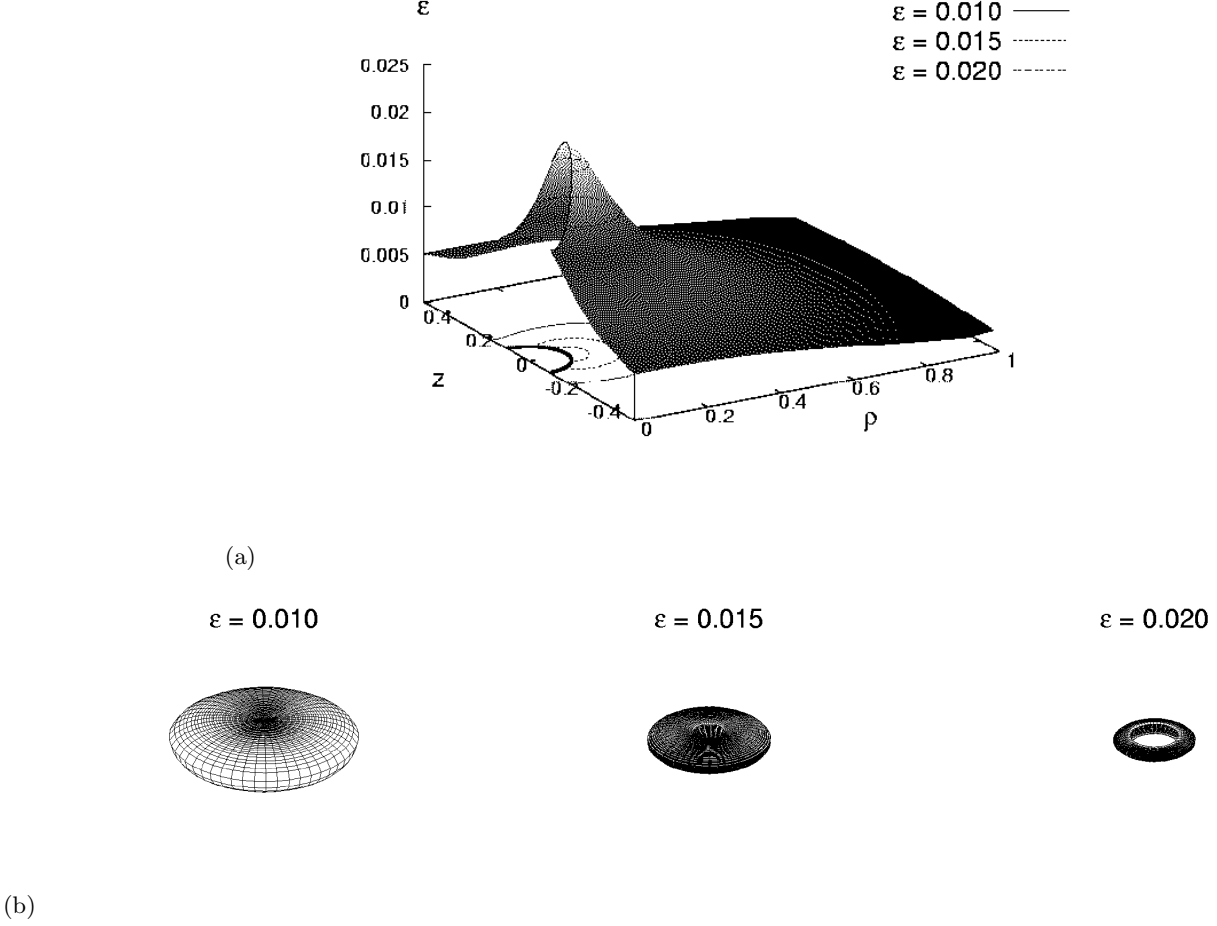


Fig. 4 (a) The energy density of the matter fields for a dyonic rotating black hole ($m = 1$, $n = 1$, $x_H = 0.1$, $\Omega = 0.5$, $\nu = 0.04$, $\alpha = 0.3$, $\beta = 0.1$, $\gamma = 0.1$). (b) Also shown are surfaces of constant energy density.

VI. CONCLUSIONS

We have considered non-perturbative globally regular and black hole solutions of EYMHD theory. These stationary axially symmetric solutions are asymptotically flat. The solutions are characterized by two integers, m and n , related to the number of monopoles and antimonopoles in the solutions, and to the magnetic charge of the monopoles, respectively [17]. The black hole solutions carry non-Abelian hair outside their regular horizon.

The globally regular solutions do not rotate, when they carry a global magnetic charge. Only solutions with no global magnetic charge can possess angular momentum, which is then quantized in terms of the electric charge [20]. The globally regular solutions satisfy a simple mass formula

$$M = \frac{D}{\gamma} + 2\nu Q .$$

The presence of electric charge enforces stationarity of the solutions, since it gives rise to an angular momentum density (except for the spherical $n = 1$ monopole).

The effective action of the globally regular solutions can be expressed in terms of the mass and the electric charge

$$S^{\text{eff}} = -(M - \nu Q) = -\left(\frac{D}{\gamma} + \nu Q\right) .$$

Based on the effective action we have shown, that the mass of stationary solutions can exhibit a “loop” close to the maximal value of the coupling constant α_{\max} , whereas the mass of static solutions can only exhibit a “spike” there (for given m , n , β , γ and ν).

Rotating EYMHD black hole solutions satisfy the zeroth and the first law of black hole mechanics. Here we have derived a non-Abelian mass formula for these black holes, which involves their global charges and their horizon properties

$$M = 2\hat{T}\hat{S} + 2\Omega J + \frac{D}{\gamma} + 2\nu Q + 2\tilde{\psi}_{\text{el}}Q(1 - \varepsilon) .$$

This mass formula differs from the EMD and EYMD mass formula Eq. (3), since the last term is present only for magnetically neutral black holes, and further an additional term is present for all electrically charged solutions, where the gauge potential parameter ν is entering together with the electric charge. When the black hole horizon size is taken to zero, the mass formula for globally regular solutions is recovered.

Whether the presence of the dilaton also allows for a new uniqueness conjecture for hairy black holes remains to be seen. Clearly, when only the mass, the angular momentum and the electric and magnetic charges are considered, the black hole solutions are not uniquely determined by these global charges.

In the numerical calculations we have only began to investigate the large parameter space for the black hole solutions. Here further investigations might reveal new phenomena, not encountered previously for non-Abelian black holes. For instance, non-Abelian counterexamples to the staticity theorem might arise as well as counterrotating black holes [6]. Also, systems of non-Abelian black holes with regular non-degenerate horizons might exist.

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- [1] W. Israel, *Commun. Math. Phys.* **8** (1968) 245;
D. C. Robinson, *Phys. Rev. Lett.* **34** (1975) 905;
P. Mazur, *J. Phys.* **A15** (1982) 3173.
 - [2] M. Heusler, *Black Hole Uniqueness Theorems*, (Cambridge University Press, 1996).
 - [3] G. W. Gibbons and K. Maeda, *Nucl. Phys.* **B298** (1988) 741;
D. Garfinkle, G. T. Horowitz and A. Strominger, *Phys. Rev.* **D43** (1991) 3140.
 - [4] R. Wald, *The first law of black hole mechanics: vol 1*, in *Directions in general relativity*, eds. B.L. Hu, M.P. Ryan, Jr., C.V. Vishveshwara and T.A. Jacobson, pp 358-366, Cambridge University Press (1993); gr-qc/9305022.
 - [5] D. Rasheed, *Nucl. Phys.* **B454** (1995) 379.
 - [6] B. Kleihaus, J. Kunz, and F. Navarro-Lérida, *Phys. Rev.* **D69** (2004) 081501.
 - [7] see e.g. M. S. Volkov and D. V. Gal'tsov, *Phys. Rept.* **319** (1999) 1;
D. V. Gal'tsov, hep-th/0112038, *Proceedings of the 16th International Conference on General Relativity and Gravitation, July 2001, Durban, South Africa*
 - [8] M. S. Volkov and D. V. Galt'sov, *Sov. J. Nucl. Phys.* **51** (1990) 747;
P. Bizon, *Phys. Rev. Lett.* **64** (1990) 2844;
H. P. Künzle and A. K. M. Masoud-ul-Alam, *J. Math. Phys.* **31** (1990) 928.
 - [9] K. Lee, V.P. Nair and E.J. Weinberg, *Phys. Rev.* **D45** (1992) 2751;
P. Breitenlohner, P. Forgacs and D. Maison, *Nucl. Phys.* **B383** (1992) 357;
P. Breitenlohner, P. Forgacs and D. Maison, *Nucl. Phys.* **B442** (1995) 126.
 - [10] B. Kleihaus and J. Kunz, *Phys. Rev. Lett.* **79** (1997) 1595;
B. Kleihaus and J. Kunz, *Phys. Rev.* **D57** (1998) 6138;
B. Kleihaus and J. Kunz, *Phys. Lett.* **B494** (2000) 130.
 - [11] B. Hartmann, B. Kleihaus, and J. Kunz, *Phys. Rev. Lett.* **86** (2001) 1422;
B. Hartmann, B. Kleihaus, and J. Kunz, *Phys. Rev.* **D65** (2002) 024027.
 - [12] B. Kleihaus and J. Kunz, *Phys. Rev. Lett.* **86** (2001) 3704.
 - [13] B. Kleihaus, J. Kunz, and F. Navarro-Lérida, *Phys. Rev.* **D66** (2002) 104001;
B. Kleihaus, J. Kunz, and F. Navarro-Lérida, *Phys. Rev.* **D69** (2004) 064028;
B. Kleihaus, J. Kunz, and F. Navarro-Lérida, *Phys. Lett.* **B599** (2004) 294.
 - [14] B. Kleihaus, J. Kunz, and F. Navarro-Lérida, *Phys. Rev. Lett.* **90** (2003) 171101.
 - [15] R. Bartnik and J. McKinnon, *Phys. Rev. Lett.* **61** (1988) 141.

- [16] G. 't Hooft, Nucl. Phys. **B79** (1974) 276;
A. M. Polyakov, JETP Lett. **20** (1974) 194.
- [17] B. Kleihaus, and J. Kunz, Phys. Rev. Lett. **85** (2000) 2430;
B. Kleihaus, J. Kunz, and Ya. Shnir, Phys. Rev. **D71** (2005) 024013;
J. Kunz, U. Neemann, and Ya. Shnir, hep-th/0703232, Phys. Rev. **D** in press.
- [18] M. S. Volkov and N. Straumann, Phys. Rev. Lett. **79** (1997) 1428;
O. Brodbeck, M. Heusler, N. Straumann and M. S. Volkov, Phys. Rev. Lett. **79** (1997) 4310.
- [19] P. Bizon and O. T. Popp, Class. Quant. Grav. **9** (1992) 193.
- [20] J. J. van der Bij and E. Radu, Int. J. Mod. Phys. **A17** (2002) 1477.
- [21] V. Paturyan, E. Radu, and D. H. Tchrakian, Phys. Lett. **B609** (2005) 360;
Y. Brihaye, B. Hartmann, and E. Radu, Phys. Rev. **D71** (2005) 085002.
- [22] B. Kleihaus, J. Kunz, and U. Neemann, Phys. Lett. **B623** (2005) 171.
- [23] P. Breitenlohner, private communication.
- [24] R. M. Wald, General Relativity (University of Chicago Press, Chicago, 1984)
- [25] L. Smarr, Phys. Rev. Lett. **30** (1973) 71;
L. Smarr, Phys. Rev. **D7** (1973) 289.
- [26] P. Forgacs and J. Gyurusi, Phys. Lett. **B366** (1996) 205;
Y. Brihaye, B. Hartmann, J. Kunz, Phys. Rev. **D65** (2002) 024019.
- [27] M. Heusler and N. Straumann, Class. Quantum Grav. **10** (1993) 1299;
Phys. Lett. **B315** (1993) 55.
- [28] G. W. Gibbons and S. W. Hawking, Phys. Rev. **D15** (1977) 2752.
- [29] P. Forgacs and N. S. Manton, Commun. Math. Phys. **72** (1980) 15.
- [30] B. Kleihaus and J. Kunz, Phys. Rev. Lett. **78** (1997) 2527;
B. Kleihaus and J. Kunz, Phys. Rev. **D57** (1998) 834.
- [31] Y. Brihaye, B. Hartmann, and J. Kunz, Phys. Lett. **B441** (1998) 77;
Y. Brihaye, B. Hartmann, J. Kunz, and N. Tell, Phys. Rev. **D60** (1999) 104016.
- [32] L. F. Abbott, and S. Deser, Phys. Lett. **116B** (1982) 259;
J. D. E. Creighton, and R. B. Mann, Phys. Rev. **D52** (1995) 4569.
- [33] W. Schönauer and R. Weiß, J. Comput. Appl. Math. **27** (1989) 279;
M. Schauder, R. Weiß and W. Schönauer, The CADSOL Program Package, Universität Karlsruhe, Interner Bericht Nr. 46/92 (1992).